Original Research Article

Non-wave solutions of the Maxwell-Einstein Equations.

ABSTRACT

This article is devoted to treating of non-wave, i.e. instanton solution for the Maxwell-Einstein equations. Equations for the field of instanton and metrics are derived. Metrics of pseudo-Euclid space which is corresponding to transition between degenerate classical vacua of problem and is connected with presence at the space infinity divergent and convergent spherical electromagnetic waves is studied. An expression of the instanton is received and it's size is found. Value of pseudo-Euclid action is calculated. It is shown that instanton violates so called "week energetic condition" which is essential for space-time singularities proving.

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Keywords: instanton, pseudo-Euclid space, classical vacuum, pseudo-Euclid action.

1. INTRODUCTION

Gravitational instantons attract attention, starting from [1]. We recall the definition. Topologically nontrivial instanton called localized solution of the classical pseudo-Euclidean field equations characterized by a finite action and connecting two different vacuums of the theory [2]. Euclidean version of the theory is introduced by replacing the Minkowski metric $g^{\mu\nu}$ $(g^{00} = 1,\ g^{ij} = -\ \delta_{ij},\ i,j=1,2,3)$ to the Euclidean metric $\delta^{\mu\nu}$. Formally, the transition from the description in Minkowski space to a description in pseudo-Euclidean space is performed by replacing the time coordinate x^0 in Minkowski space to coordinate $y^0 = ix^0$ in pseudo-Euclidean space, while introducing a pseudo-Euclidean action Λ, associating it with the action in Minkowski space S by expression $\Lambda = iS$, $i = (-1)^{½}$.

Instantons of classical field equations in Minkowski space describe in the semiclassical approximation quantum tunneling process between degenerate classical states located near a variety of vacuums. In the theory of Maxwell-Einstein (M-E) equations. These degenerate states are states in which there is convergent (divergent) electromagnetic wave at spatial infinity, which represent the two degenerate vacuums of the theory. As was shown in [3] classical transition between these states is impossible. Indeed, if we consider the vacuum in which there is a convergent spherical electromagnetic wave (SEMW), then taking into account the curvature of space-time due to the waves almost all the rays corresponding to small portions of the wave front will capture by curvature of metrics and do not give a contribution to the outgoing wave[1]. Therefore, the role of the instanton of the M-E equations

[1] In other words, convergent wave is not focused to a point, or, in mathematical terms corresponding map is not homotopic to zero [4].

is extremely important to description such an intuitive and "simple" phenomenon, which seems to be the process of transformation a convergent SEMW to a divergent one. Another important application of the instanton of the M-E equations is to develop a physical theory of electromagnetic resonators, which eliminates the unphysical singularities of fields, for example, in a spherical cavity [5]. And at last, it is possible to be an important application of the theory developed include cosmology, because the process of transformation of a convergent to a divergent SEMW is one of the main processes in the universe.

2. BASIC EQUATIONS

As the initial equations we choose the Einstein's gravitational equations and the equations of the electromagnetic field in vacuum (Maxwell's equations) associated with each other [6]:

$$R_{ik}-\frac{1}{2}g_{ik}R=\frac{8\pi K}{c^4}T_{ik};F^{ik}_{,k}+\Gamma^{l}_{kl}F^{ik}=0 \qquad (1)$$

Here R – trace of Ricci's tensor R^i_k: $R = R^i_i$, g_{ik} – metric tensor; T_{ik} and F^{ik} – tensor of energy-momentum and electromagnetic one; Γ^i_{kl} – Christoffel's symbols; c – light speed in vacuum, K – gravitation constant; indices i, k, l take values 0, 1, 2, 3; repeated indices mean summation; comma means usual, i.e. non-covariant derivative[7]. Let us find a solution of (1) which corresponds to existence of spherical light wave at $r \to \infty$. For this we use an expression for interval just as in well-known Schwarzschild problem[7]:

$$ds^2=e^{\nu}c^2dt^2-e^{\lambda}dr^2-r^2(d\theta^2+\sin^2\theta\cdot d\varphi^2) \qquad (2)$$

$\nu = \nu(t, r, \theta)$, $\lambda = \lambda(t, r, \theta)$; $x^0 = ct$, t –time; $x^1 = r$, $x^2 = \theta$, $x^3 = \varphi$ – spherical co-ordinates. SEMW is characterized by frequency ω and kinetic moment vector $\boldsymbol{M}$. Let us choose z - axis of the co-ordinate system in direction perpendicular to $\boldsymbol{M}$. It simplifies a treating because the dependence of azimuth angle φ in (1) may be omitted.

Second equation in (1) may be transformed to [7]:

$$\frac{\partial}{\partial x^{\beta}}\left(\sqrt{-g}F^{\alpha\beta}\right)-\frac{\partial}{\partial x^{0}}\left(\sqrt{-g}F^{0\alpha}\right)=0$$
$$\frac{\partial}{\partial x^{\beta}}\left(\sqrt{-g}F^{0\beta}\right)=0,$$
$$\alpha,\beta=1,2,3;-g=e^{\frac{\lambda+\nu}{2}}r^2\sin^2\theta \qquad (3)$$

where *α =1, 2* correspond for the SEMW of TM - type, and *α=3* – for SEMW of TE-type. Below we restrict ourselves with the case of TM - type[2], for which nonzero components of vector-potential and electromagnetic tensor are only *A1, A2* and:

$$F_{01}=\frac{\partial A_1}{\partial x^0},F_{12}=\frac{\partial A_2}{\partial x^1}-\frac{\partial A_1}{\partial x^2},F_{02}=\frac{\partial A_2}{\partial x^0} \tag{4}$$

We use the Hamilton calibration $A_0 = 0$. For variables' separation we assume an additional condition: $\lambda = \alpha(r, t)+\beta(\theta)$, $\nu = -\alpha(r, t)+\beta(\theta)$. Substitution (4) into (3) gives us two equations for the components A_1 and A_2:

$$e^{-\alpha(r,t)}\sin\theta\frac{\partial}{c\partial r}\left[r^2e^{-\beta(\theta)}\frac{\partial A_1}{\partial t}\right]+\frac{\partial}{c\partial\theta}\left[\sin\theta\frac{\partial A_2}{\partial t}\right]=0$$
$$\frac{\partial}{\partial\theta}\left[\sin\theta\left(\frac{\partial A_1}{\partial\theta}-\frac{\partial A_2}{\partial r}\right)\right]-e^{\alpha(r,t)}\frac{r^2}{c^2}\sin\theta\frac{\partial}{\partial t}\left[e^{-\beta(\theta)}\frac{\partial A_1}{\partial t}\right]=0 \tag{5}$$

If we take a derivative of the first equation in (5) on *r* and second – on *ct,* then we exclude A_2 from the equations. Representing $F_{01} = \Psi(r,t)\cdot\Phi(\theta)$, we receive equations for $\Psi(r,t)$ and $\Phi(\theta)$:

$$\frac{d^2\Phi}{d\theta^2}+ctg\theta\frac{d\Phi}{d\theta}+l(l+1)e^{-\beta(\theta)}\Phi=0$$
$$e^{-\alpha(r,t)}\frac{\partial^2}{\partial r^2}\left(r^2\Psi\right)-e^{\alpha(r,t)}\frac{\partial^2}{c^2\partial t^2}(r^2\Psi)-l(l+1)\Psi=0 \tag{6}$$

It leads from (6) that for $\beta \to 0$ $\Phi(\cos\theta) = P_l(\cos\theta)$, where $P_l(\cos\theta)$ – Legendre polinomial, and *l* is nonnegative integer[8].

Energy-momentum tensor's components T_k^i may be expressed by the components of metric tensor g_k^i with the help of Einstein's equation of gravity [7]:

$$\frac{8\pi K}{c^4}T_k^i=R_k^i-\frac{1}{2}\delta_k^iR \tag{7}$$

where δ_k^i – unit 4-tensor, and *R* – is a trace of tensor R^i_k. The details of calculations one can find in [7], for example. Besides Christoffel's symbols presented in[7], we need some additional ones; a symbol ˜ means a derivative by the angle *θ* :

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[2] A solution for the SEMW of *TE* – type does not need separate treating because Maxwell-Einstein equations are invariant with the transformation ***E* → -*H*, *H* → *E*,** (so as Maxwell ones) due to invariance of energy-momentum tensor T_{ik}

$$\Gamma^{1}_{12} = \Gamma^{1}_{21} = \frac{\tilde{\lambda}}{2}, \Gamma^{0}_{02} = \Gamma^{0}_{20} = \frac{\tilde{\nu}}{2},$$
$$\Gamma^{2}_{11} = \frac{\tilde{\lambda} e^{\lambda}}{2r^{2}}, \Gamma^{2}_{00} = \frac{\tilde{\nu} e^{\nu}}{2r^{2}}$$

A result looks as follows:

$$\frac{8\pi K}{c^4} T^{0}_{0} = -e^{-\alpha-\beta}\left(\frac{1}{r^2} - \frac{\alpha'}{r}\right) + \frac{1}{r^2} + \frac{\tilde{\beta}}{2r^2}\left(2\tilde{\beta} + 1\right);$$
$$\frac{8\pi K}{c^4} T^{1}_{1} = -e^{-\alpha-\beta}\left(\frac{1}{r^2} - \frac{\alpha'}{r}\right) + \frac{1}{r^2} -$$
$$\frac{1}{r^2}\left[2\tilde{\tilde{\beta}} + 2\left(\tilde{\beta}\right)^2 + 2\tilde{\beta} ctg\,\theta - \frac{\tilde{\beta}}{2}\right];$$
$$\frac{8\pi K}{c^4} T^{2}_{2} = \frac{1}{2} e^{-\alpha-\beta}\left[\alpha'' - \left(\alpha'\right)^2 + \frac{2\alpha'}{r}\right] +$$
$$\frac{1}{2} e^{\alpha-\beta}\left(\ddot{\alpha} + \dot{\alpha}^2\right) + \frac{1}{2r^2}\left[\tilde{\tilde{\beta}} + \left(\tilde{\beta}\right)^2 - \tilde{\beta} ctg\,\theta - \frac{\tilde{\beta}}{2}\right];$$
$$\frac{8\pi K}{c^4} T^{3}_{3} = \frac{1}{2} e^{-\alpha-\beta}\left[\alpha'' - \left(\alpha'\right)^2 + \frac{2\alpha'}{r}\right] +$$
$$\frac{1}{2} e^{\alpha-\beta}\left(\ddot{\alpha} + \dot{\alpha}^2\right) - \frac{1}{2r^2}\left(\tilde{\tilde{\beta}} - \tilde{\beta} ctg\,\theta + \frac{\tilde{\beta}}{2}\right);$$
$$\frac{8\pi K}{c^4} T^{1}_{0} = -e^{-\alpha-\beta}\frac{\dot{\alpha}}{r}; \frac{8\pi K}{c^4} T^{2}_{0} = 0;$$

$$\frac{8\pi K}{c^4} T^{2}_{1} = -\frac{2\tilde{\beta}}{r^3} \qquad (8)$$

Using these expressions we rewrite final equations in the form:

$$\frac{8\pi K}{c^4}T_0^0 = \frac{2K}{c^4}\left\{\frac{1}{2}e^{-2\beta}\Psi^2 + \frac{1}{[l(l+1)]^2}\left[\frac{1}{2r^2}e^{-\alpha-\beta}\left(\frac{\partial}{\partial r}r^2\Psi\right)^2 + \frac{r^2}{2}e^{\alpha-\beta}\left(\frac{\partial \Psi}{\partial x^0}\right)^2\right]\right\}\Phi^2$$

$$\frac{8\pi K}{c^4}T_1^1 = \frac{2K}{c^4}\left\{\frac{1}{2}e^{-2\beta}\Psi^2 - \frac{1}{[l(l+1)]^2}\left[\frac{1}{2r^2}e^{-\alpha-\beta}\left(\frac{\partial}{\partial r}r^2\Psi\right)^2 + \frac{r^2}{2}e^{\alpha-\beta}\left(\frac{\partial \Psi}{\partial x^0}\right)^2\right]\right\}\Phi^2$$

$$\frac{8\pi K}{c^4}T_2^2 = \frac{2K}{c^4}\left\{-\frac{1}{2}e^{-2\beta}\Psi^2 + \frac{1}{[l(l+1)]^2}\left[\frac{1}{2r^2}e^{-\alpha-\beta}\left(\frac{\partial}{\partial r}r^2\Psi\right)^2 - \frac{r^2}{2}e^{\alpha-\beta}\left(\frac{\partial \Psi}{\partial x^0}\right)^2\right]\right\}\Phi^2$$

$$\frac{8\pi K}{c^4}T_3^3 = \frac{2K}{c^4}\left\{-\frac{1}{2}e^{-2\beta}\Psi^2 + \frac{1}{[l(l+1)]^2}\left[-\frac{1}{2r^2}e^{-\alpha-\beta}\left(\frac{\partial}{\partial r}r^2\Psi\right)^2 + \frac{r^2}{2}e^{\alpha-\beta}\left(\frac{\partial \Psi}{\partial x^0}\right)^2\right]\right\}\Phi^2$$

$$\frac{8\pi K}{c^4}T_0^1 = -\frac{2K}{c^4}\frac{e^{-\alpha-\beta}}{[l(l+1)]^2}\frac{\partial \Psi}{\partial x^0}\left(\frac{\partial}{\partial r}r^2\Psi\right)\Phi^2; \frac{8\pi K}{c^4}T_0^2 = \frac{2K}{c^4}\frac{e^{-\beta}}{l(l+1)}\Psi\frac{\partial \Psi}{\partial x^0}\cdot\Phi^2$$

$$\frac{8\pi K}{c^4}T_1^2 = \frac{2K}{c^4}\frac{e^{-\beta}}{l(l+1)}\frac{\Psi}{r^2}\frac{\partial}{\partial r}\left(r^2\Psi\right)\cdot\Phi^2 \qquad (9)$$

In accordance with [6] right sides of equations (9) is averaged over the angle *θ*. In addition, they are for the wave solutions is averaged over time [6]. For the non-wave-solutions making procedure of averaging over time has no meaning. Consider first the last three equations: for the T_0^1, T_0^2 and T_1^2. Note, that their right hand sides, except the equation for the T_0^1 are of the order of value $\sim r_s^2/r^2 << 1$, where $r_s^2 = K < f >^2/2c^4$, $< f >$- solution's order of value, $f = r^2\Psi$. Therefore, at distances of the order of the wavelength of light right hand sides of equations can be omitted[3]. This is consistent with the equations for T_0^2 and T_1^2 (5), if $\widetilde{\beta} = 0$ .The equation for T_0^1 in (9) we will use to find *α*.

3. TREATMENT THE EQUATIONS

Subtracting the second equation in (9) from the first one and equating the result with the same operation which is done with (8) we receive:

$$\frac{2K}{c^4}\left[e^{-\alpha}\left(\frac{\partial}{\partial r}r^2\Psi\right)^2 + e^{\alpha}\left(\frac{\partial}{\partial x^0}r^2\Psi\right)^2\right] = \frac{e^{\beta}[l(l+1)]^2}{\Phi^2}\left[\widetilde{\widetilde{\beta}} + 2\widetilde{\beta}^2 + \widetilde{\beta}\left(ctg\theta + \frac{1}{4}\right)\right] = A \qquad (10)$$

[3] See details in [3].

A is a constant. The solutions of (6), corresponding to the equation (10) one can treat in pseudo-Eucleadean space which metric follows from the Minkowski space's metric with substitution time co-ordinate x^0 to "time" co-ordinate $-iy^0$ in pseudo-Euclidean space. At the same time one can introduce pseudo-Euclead action Λ, which is connected with the action *S* in Minkowski space as follows $\Lambda = iS$, $i = (-1)^{½}$. It is known [2] that localized solutions of Euclidean field equations with finite Euclidean action are instantons. An instantons of classical field equations in Minkowski space describe in quasi-classical limit tunneling between degenerate classical states, which contain convergent and divergent SEMW. This procedure turns second hyperbolic equation in (6) to the elliptic one. If one suppose its finiteness at $r \to \infty$ then he receives a condition $A = 0$ from the equation (10). This provides second equation (6) looks as follows:

$$e^{\alpha} = \pm \frac{\partial f}{\partial r}\left(\frac{\partial f}{\partial y^0}\right)^{-1};$$

$$f'' + \left(\frac{f'}{\dot f}\right)^2 \ddot f \mp \frac{l(l+1)}{r^2}\frac{f'}{\dot f} f = 0;$$

$$f = r^2\Psi \tag{11}$$

Here prime still means a derivative on $x^1 = r$, and point – on $y^0 = c\tau$.

Signs ± hereafter correspond to different vacuums of the theory, located at $\tau \to \pm\infty$. Using (11) we can rewrite the equations for instantons (8) and (9) in the form:

<u>Einstein equations:</u>

$$-e^{-\alpha}\left(\frac{1}{r^2} - \frac{\alpha'}{r}\right) + \frac{1}{r^2} = \frac{K}{c^4}\frac{f^2}{r^4}$$

$$\frac{1}{2}e^{-\alpha}\left[\alpha'' - (\alpha')^2 + \frac{2\alpha'}{r}\right] - \frac{1}{2}e^{\alpha}\left(\ddot\alpha + \dot\alpha^2\right) = -\frac{2K}{c^4}\left\{\frac{f^2}{2r^4} + \frac{1}{[l(l+1)]^2}\frac{e^{\alpha}}{r^2}\left(\dot f\right)^2\right\} \tag{12}$$

Point here and below means derivative y^0.

<u>Maxwell equations:</u>

$$f'' + e^{2\alpha}\ddot f - e^{\alpha}\frac{l(l+1)}{r^2} f = 0 \tag{13}$$

Consider the equation (12). They are compatible if the condition is true:

$$e^{\alpha}\left(\ddot{\alpha}+\dot{\alpha}^{2}\right)\mp\frac{2\dot{\alpha}}{r}=\frac{K}{c^{4}}\frac{1}{r^{3}}\frac{df^{2}}{dr} \quad (14)$$

Condition (14) allows us to rewrite equations (12) in the form:

$$-e^{-\alpha}\left(\frac{1}{r^{2}}-\frac{\alpha'}{r}\right)+\frac{1}{r^{2}}=\frac{K}{c^{4}}\frac{f^{2}}{r^{4}}$$
$$e^{-\alpha}\left[\alpha''-(\alpha')^{2}+\frac{2\alpha'}{r}\right]=-\frac{2K}{c^{4}}\left[\frac{f^{2}}{2r^{4}}+\frac{1}{r^{3}}\frac{df^{2}}{dr}\right] \quad (15)$$

Solve these equations is very difficult even numerically. Therefore, we are interested mainly in the asymptotic behavior of their solutions at distances of a wavelength of light (or more). Note that their right sides are of the order ~ $r_s^2/r^2 \ll 1$ and they can be omitted with the adopted accuracy. Then Einstein's equations reduce to a single equation. It solution is consistent with the metric in a space free of matter and is well known

$$e^{\alpha}\approx 1+\frac{const}{r} \quad (16)$$

where the value of the constant *const* is to be determined. Furthermore, this asymptotic equation (13) and (14) have autoscale solutions, depending on $z = c\tau/r$. For such solutions instead of (3.7) we obtain the equation

$$\sigma''\mp 2\frac{\sigma'}{\sigma}=0,\sigma=e^{\alpha},\sigma'=\frac{d\sigma}{dz} \quad (17)$$

Equation (17) is easily integrated and leads to the expression $Ei(\alpha)=\pm 2z$, where *Ei*– integral exponent. Using the well-known expansion [9]:

$$Ei(\alpha)=\ln|\alpha|+\sum_{k=1}^{\infty}\frac{\alpha^{k}}{k\cdot k!}$$

we can get the expression for the metric for large values of *z*:

$$e^{\alpha}\approx 1+e^{-2|z|} \quad (18)$$

Formula (18) describes the transition between the vacuum states with flat metric corresponding to the presence of at $z \to -\infty$ convergent SEMW, and at $z \to +\infty$ divergent one. This transition is localized to τ, the localization region has a size ~ r/c. Einstein's equations are also satisfied because "time" τ does not appear in them, and the equation (17) has a solution that can be represented for small *z* (large *r*) as a series

$$\sigma = e^{\alpha} = 1 + \mu z - \mu z^2 + \frac{\mu(\mu+1)}{3} z^3 + \ldots \approx 1 + \frac{\mu\tau}{r}, \mu = \sigma'(0) \qquad (19).$$

4. PSEUDO-EUCLID ACTION

Let us calculate the action in curved space-time [7]

$$S_f = -\frac{1}{16\pi c}\int F_{ik}F^{ik}\sqrt{-g}\,d\Omega;$$
$$d\Omega = dx^0 dx^1 dx^2 dx^3;$$
$$\sqrt{-g} = r^2 e^{\beta(\theta)} \sin\theta \qquad (20)$$

Turning to action in pseudo-Euclidean space $\Lambda = iS_f$, $dx^0 = -icd\tau$, and using (11) and the normalization $\Phi(\theta)$ [8], we receive (if $\beta = 0$)

$$\Lambda = \frac{1}{4c^2}\int\left\{\left(\frac{\partial A_r}{\partial\tau}\right)^2 \pm \frac{2r^2}{c\left[l(l+1)\right]^2}\frac{\partial^2 A_r}{\partial r\partial\tau}\frac{\partial^2 A_r}{\partial\tau^2}\right\}r^2 dr d\tau \qquad (21)$$

In (21) also is taken into account that A_θ can be expressed through A_r [6]. Let us treat extremes of Λ. For this we calculate the variation Λ on A_r provided condition $\delta A_r = 0$ on the boundaries of integration and equate it to zero. As a result, we obtain:

$$\delta\Lambda = \frac{1}{2c^2}\int \delta A_r \frac{\partial^2}{\partial\tau^2}\left\{A_r \pm \frac{r^2}{c\left[l(l+1)\right]}\frac{\partial^2 A_r}{\partial r\partial\tau}\right\}r^2 dr d\tau = 0 \qquad (22)$$

Because of the arbitrariness δA_r integrand in (22) is equal zero, which gives the equation of the instanton:

$$\frac{\partial A_r}{\partial\tau} \pm \frac{r^2}{c\left[l(l+1)\right]^2}\frac{\partial^3 A_r}{\partial r\partial\tau^2} = 0 \qquad (23)$$

It reduces to the equation:

$$z\Upsilon'' \mp \left[l(l+1)\right]^2 \Upsilon = 0$$
$$\Upsilon = \frac{\partial A_r}{\partial\tau}, z = \frac{c\tau}{r} \qquad (24)$$

The solution of this equation has the form:

$$\Upsilon(z) = \sqrt{z} Z_1 \left(2l(l+1)\sqrt{\mp z} \right) \tag{25}$$

Z_1- cylindrical function. Below, however, we will use another solution of equation (24), since the solution *Y(z)* does not have finite action. Calculating the pseudo-action for the instanton $A_r^I(r,\tau)$, we receive

$$\Lambda\left(A_r^I\right) = \frac{3}{4c^2} \int \left(\frac{\partial A_r^I}{\partial \tau} \right)^2 r^2 dr d\tau > 0 \tag{26}$$

Action must be calculated for a classical trajectory begins and ends ($\tau \to \pm\infty$) in the region where the space-time is not curved, i.e. at $r \to \infty$, where the field of SEMW tends to zero. Among the set of solutions of equation (23) satisfying this condition, we choose the solution:

$$\frac{\partial A_r^I}{\partial \tau} = cE \exp\left\{ \mp \omega\tau - \frac{c}{\omega} \left[l(l+1) \right]^2 \frac{1}{r} \right\} \tag{27}$$

where *E* - is a constant with the dimension of the electric field. Its value is related to the so-called topological charge of the instanton

$$Q = \int_{-\infty}^{\infty} \frac{\partial A_r^I(r=\infty,\tau)}{\partial \tau} d\tau = 2E\frac{c}{\omega} \tag{28}$$

In evaluating the integral in (28) we use the expression (27)[4]. An important feature of the solution (27) is that it decreases the field $r \to 0$, which is consistent with the tunneling nature of the instanton.

In calculating the pseudo-Euclidean action for the solution (27) to avoid divergence of the integral in (26) we cut off the integral over *dr* in the upper limit at the distance r_0, having a sense of the size of the instanton, which will be defined below. With this in mind, the result of calculations (3.19) looks as follows:

$$\Lambda(A_r^I) = 6\frac{E^2 r_c^3}{\omega} \left[l(l+1) \right]^3 K(l)$$

$$K(l) = \int_{x_0}^{\infty} \frac{e^{-x}}{x^4} dx, x_0 = 2\frac{r_c}{r_0} l(l+1) \tag{29}$$

Recall that the action $\Lambda(A_r^I)$ determines the probability *w* of conversion of convergent SEMW in divergent one: $w \sim \exp\left(\Lambda(A_r^I)\right)/\hbar$, $\hbar = h/2\pi$, *h* – is the Plank constant [2].

[4] Typically, the integral in (28) is normalized to the right side of the equation that leads to values of Q = 1 (iinstanton) and Q = -1 (anti-instanton) [10].

The value r_0 we will find from the condition of matching metrics inside and outside the instanton. In the outer region metric is given by [3].

$$e^{\alpha}_{out} = \left[1 - \frac{r_c}{r} + \left(\frac{r_s}{r}\right)^2\right]^{-1} \tag{30}$$

Metric in the inner region is found from formula (11), where we substitute the solution (27), given that $f = ir^2/c\partial A_r/\partial\tau$

$$e^{\alpha}_{in} = \frac{2r_c}{r}\frac{1}{l(l+1)} + \left(\frac{r_c}{r}\right)^2 \tag{31}$$

Given that $r_s << r_c$ and living in equation $e^{\alpha}{}_{in} = e^{\alpha}{}_{out}$ most significant members we get (up to terms ~ $[l(l+1)]^{-1}$)

$$r_0 = r_c\left[1 - \frac{2}{3l(l+1)}\right]^{-1} \tag{32}$$

Let us calculate the amount of $R^{00} = \frac{8\pi K}{c^4}T^{00}$, using (8) for the metric (31) where $T^{00} = W$ – field energy density of the instanton.

$$R^{00} = g^{00}R^0_0 = e^{\alpha}\left[\frac{1}{r^2} + e^{-\alpha}\left(\frac{\alpha'}{r} - \frac{1}{r^2}\right)\right]; e^{\alpha} = e^{\alpha}_{in} \tag{32}$$

The result of calculation is shown in Fig. 1

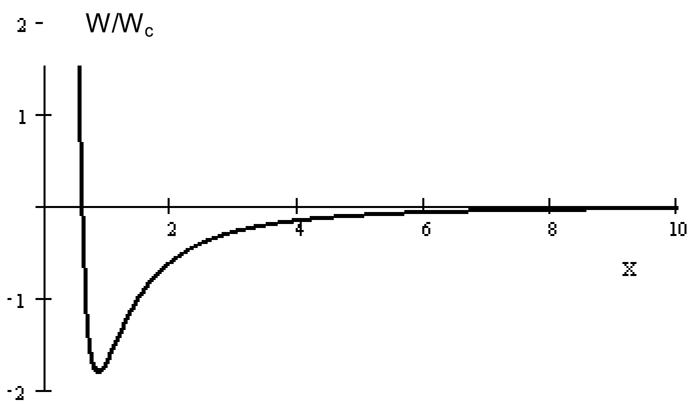

Fig. 1. Energy density of the instanton W(*x*); $W_c = \frac{c^4}{8\pi K r_c^2}$, $x = r/r_c$, $l = 3$. Qualitative behavior of W (x) is stood the same for any l.

The calculation showed that near the boundary of the instanton and waves, i.e. for $x \approx x_0$ $(r \approx r_0)$ energy density is negative. The calculation showed that the magnitude $R_{00} < 0$ in a sufficiently large vicinity of $x = x_0$ for all l. The latter circumstance is essentially for the proof of the presence (or rather lack of it) of singularities, which, as is known, is based on the fact $R_{\alpha\beta}\, \xi^\alpha\, \xi^\beta > 0$, where ξ – is any non space-like 4 – vector[5] [11]. Absence of singularities associated with horizons of the metric (30), can be seen both from the expression (31 and from Fig. 1. The only fatal singularity is a singularity at $x = 0$, where $W(x) \sim x^{-4}$.

5. PROBLEM OF INSTANTONS FROM ENERGETIC POINT OF VIEW

During the propagation of SEMW part of its energy converts into other energy forms, such as energy of the gravitational waves. This issue was outside the scope of the work [see 3, 5, 6].

In the literature there are different points of view on the question on interaction of EMW and gravitational waves. In [12] argues that the processes of transformation of the two photons in the graviton (and back) are prohibited by the conservation laws. At the same time, Wheeler did not rule out such a possibility [13]. In [14], these processes are considered without any discussion. These differences can be overcome, if we consider the photon-graviton processes in the presence of a static gravitational field created by SEMW, which removes the restrictions imposed by the conservation laws[6]. Leaving this issue for further discussion, make the following remark. Consider the relation (3.4)

$$e^\alpha = \pm c \frac{\partial f}{\partial r}\left(\frac{\partial f}{\partial \tau}\right)^{-1}; f = r^2\Psi, \tau = i\frac{x^0}{c} \qquad (33)$$

for instantons, and a similar relation for waves

$$e^\alpha = \pm \frac{\partial f}{\partial r}\left(\frac{\partial f}{\partial x^0}\right)^{-1}; f = r^2\Psi \qquad (33a)$$

binding metrics and fields of instanton or electromagnetic waves. For definiteness we take in (33) and (33a) "+" sign. Using Maxwell's equations in curved space-time [7] give them to the form, respectively

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[5] For proving of $R_{\alpha\beta}\, \xi^\alpha\, \xi^\beta < 0$ one can take, for example, ξ *(1, 0, 0, 0).*

[6] Like that the diagrams with three free ends become possible in quantum electrodynamics [15]

$$e^{-\frac{\alpha}{2}} = 1 - i\int_{r}^{\infty} \frac{(rot\vec{H})_r}{\vec{E}_r}\,dr$$

$$e^{-\frac{\alpha}{2}} = 1 - \int_{r}^{\infty} \frac{(rot\vec{H})_r}{\vec{E}_r}\,dr \qquad (34)$$

$\vec{E}, \vec{H}$ - electric and magnetic fields of the instanton or SEMW. The magnitude of the integrand is proportional to the conductivity (the role of the current density plays displacement current density), the value for which is real for wave and imaginary for instanton. The first means that the energy irreversibly converts from SEMW in some other form, which is most likely connected with gravitational waves. The second indicates the reversible transfer of energy from the electromagnetic wave to the instanton, with subsequent return to the SEMW.

A question of interest is that, at what stage of the study was the neglect of gravitational waves, and what role they play in the problem. If we argue by analogy with the problem of the gravitational collapse of a non-spherical body, it can be assumed that the emission of gravitational waves will accompany the propagation of a spherical electromagnetic wave with a nonzero *l*, that, in the end of ends allow to speak about a spherically symmetric metric for *l* ≠ 0. Thus, used in this work, as well as in [3, 5, 6], averaging tensor T_i^k (9) on the angle θ as a consequence led to the fact that gravitational waves have been left out of consideration.

6. DISCUSSION

This article is devoted to treating the role of instantons in considering the dynamics of spherical electromagnetic waves by means of Maxwell-Einstein equations. Thanks instantons convergent wave can be transformed into a divergent one what allow transmission of information from the past to the future. This article discusses the two different solutions of it – an auto-scaled one depending on $z = c\tau / r$ (25) which does not have a finite Euclidean action, and the solution (27) with the finite action $\Lambda(A_r^l)$ (29). Feature of the first solution is that in a world where it could be realized, the past is separated from the future with an infinite barrier, i.e. there is no flow of time in this world. The second solution is more consistent with the state of affairs in the real world - past goes to the future with some finite probability .

The result obtained above, consisting in violation by instantons of the so-called "weak energy condition" $T_{\alpha\beta}\,\xi^\alpha\,\xi^\beta > 0$, where ξ – is any non space-like 4 – vector is important in research of the space-time singularities [12].

Note that most of the work on gravitational instantons available on the resource [16], are devoted to the classification of instanton solutions of Maxwell-Einstein in multidimensional Riemannian manifolds and their applications to the physics of black holes.

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